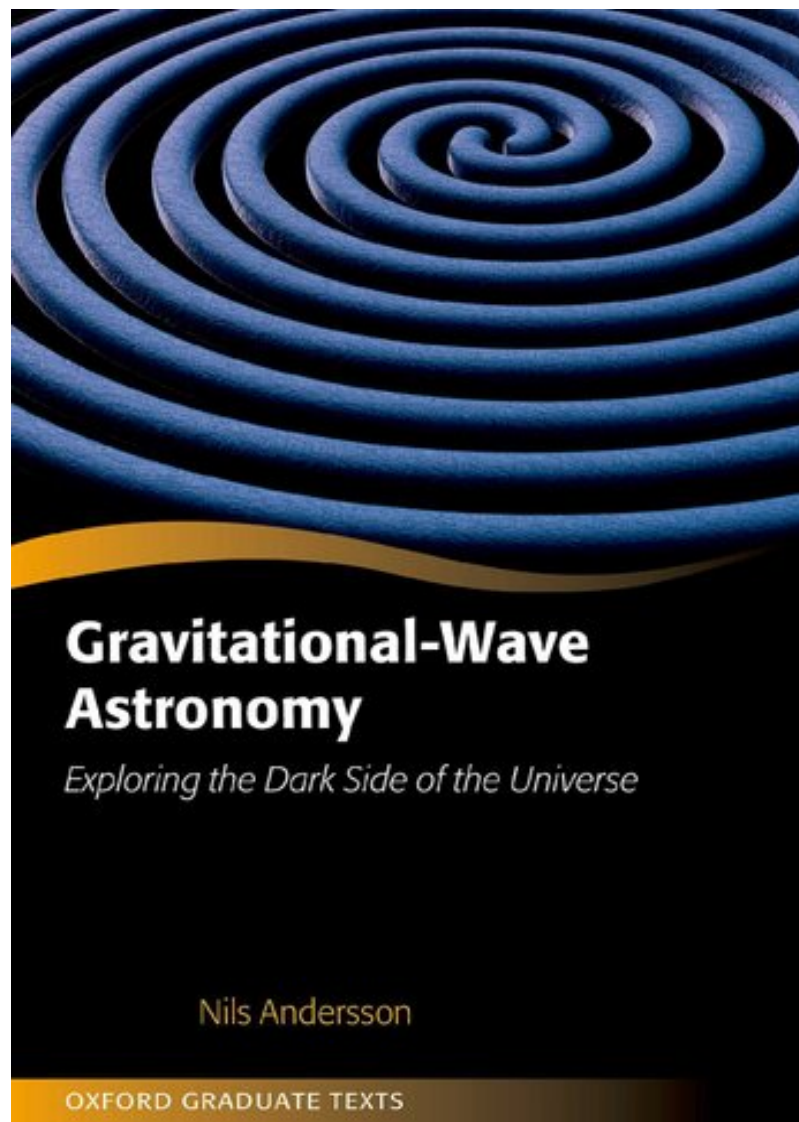


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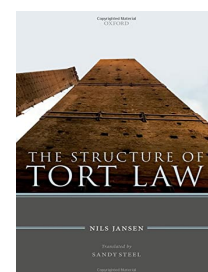
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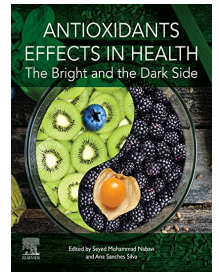
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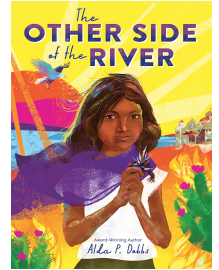
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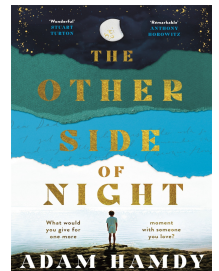
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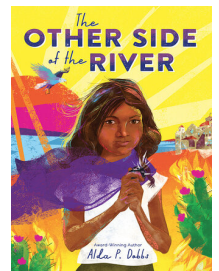
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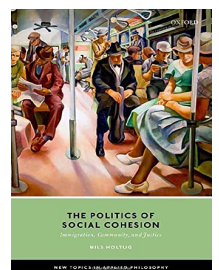
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Gravitational-Wave Astronomy

Exploring the Dark Side of the Universe

Nils Andersson

OXFORD GRADUATE TEXTS

GRAVITATIONAL-WAVE ASTRONOMY

Gravitational-Wave Astronomy

Exploring the Dark Side of the Universe

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Preface

Whenever you are writing a book, people are bound to ask: ‘What kind of book is it? Who is it for?’ These questions are reasonable, but the answers may not be that obvious. You may, for example, have embarked on the project simply because it seemed like a good idea at the time. So, with this in mind, what kind of a book is this? Having lived with it for longer than I care to figure out, I still find it difficult to give a clear answer. It is much easier to explain what it is not. This book is not an exhaustive review of gravitational-wave astronomy. At least not in the sense that it provides a ‘complete’ reference list and a detailed account of the historical developments of the ideas and the scope of the field. It is much more ‘subjective’ than that. This may be frustrating to colleagues that have contributed to the developments over the last several decades, but the reality is that I had to make choices. It was simply not manageable to peek into (and report back on) every nook and cranny, no matter how fascinating this might have been. Instead, I have tried to provide an entry point to the vast (and rapidly growing!) literature on the different aspects of gravitational waves and related astrophysics.

In essence, I have tried to build a bridge across different areas of physics that have fascinated me for a long time. On the one hand, we have gravity—with Einstein’s warped spacetime providing an astonishing example of what the human mind is capable of. On the other hand, there is the extraordinary range of astrophysics and cosmology that comes into play when we try to understand the gravitational-wave sky. And finally, we need to consider the sublime technology that was developed to catch these faint whispers from the distant Universe. This book maps out a journey through this complex landscape—introducing a combination of overlapping areas of research, many of which require their separate books for a fair treatment. The different chapters (especially in the second part) are intended to narrow the gap between a basic understanding and current research. An important part of this involves introducing the relevant language—making the involved concepts less ‘mysterious’.

The book is intended to work as a platform, sufficiently low that anyone with an interest in gravitational waves can scramble onto it, but at the same time high enough that it connects with current research—and exciting discoveries that are happening right now. It may only be an introduction, but I think it has potential... If you are an astronomer and you want a basic understanding of this new window to the Universe, including a brief (relatively self-contained) glimpse at Einstein’s theory, then this book may work for you. Similarly, if you spend most of your time analysing data from gravitational-wave detectors and you would like a better picture of what you are looking for (and perhaps why theorists find it so difficult to make firm predictions) then other parts of the book could work for you. Finally, there is a connection to nuclear physics—which is natural, since gravitational-wave signals from neutron stars may help constrain our ideas

for matter at extreme densities. Relevant aspects are addressed at various places in the book, which may help nuclear and particle physicists appreciate how their work fits into the bigger picture. Whichever direction you are coming from, and regardless of where you are going, this book may be of interest to you.

In terms of teaching, the scope of the book is likely too vast for a single undergraduate or masters-level course. But the material is flexible. The first part introduces the key ideas, following a general overview chapter and including a brief reminder of Einstein's theory. This part can be taught as a (fairly) self-contained undergraduate one semester course. In fact, the material is based on a course we have had on the books for over a decade. So I know it works. Depending on the background and interest of the students, I would select topics from the second (much longer) part of the book to connect with the actual state of the art. The chapters are written to work as 'set pieces' with core material that can be adapted to specific lectures and additional material that provide context and depth. At least that's the way I like to think about it. Some of the chapters have been road-tested at summer schools and other events so I am confident they work. The one thing that is missing in terms of teaching material is exercises. However, it is quite easy to identify steps that need filling in and to come up with questions that go beyond the material, so this should not be a major issue.

Before we embark on the journey, it is useful to make a few comments on notation and conventions. Throughout the book I have chosen to work with a spacetime metric with signature $+2$. There is one exception: The discussion of the Newman–Penrose formalism used to discuss the dynamics of spinning black holes. I have adopted the convention that spacetime indices are given by letters from the beginning of the alphabet, a, b, c, \dots , while spatial indices start with i, j, k, \dots . Many text books use Greek letters for the former. Repeated indices (spacetime or spatial) indicate summation.

With these formalities out of the way, let's get started.

Contents

1	Opening the window	1
1.1	The beginning	1
1.2	A new kind of astronomy	3
1.3	Audio not video	6
1.4	On the back of an envelope	7
1.5	Binary inspiral and merger	10
1.6	Supernovae	14
1.7	Spinning neutron stars	15
1.8	Fundamental physics	18
1.9	Many different messengers	18
1.10	The golden binary	19
 Part 1 From theory to experiment		
2	A brief survey of general relativity	25
2.1	A simple thought experiment	27
2.2	The tidal tensor	28
2.3	Introducing the metric	31
2.4	The four-velocity	34
2.5	The covariant derivative	39
2.6	The geodesic equation	41
2.7	Curvature	43
2.8	A little bit of matter	45
2.9	Geodesic deviation and Einstein's equations	47
3	Gravitational waves	51
3.1	Weak waves in an otherwise flat spacetime	52
3.2	Effect on matter	54
3.3	The wave equation	56
3.4	Transverse-traceless (TT) gauge	58
3.5	The quadrupole formula	61
3.6	The energy carried by gravitational waves	64
3.7	The radiation reaction force	67
3.8	The radiated angular momentum	70
3.9	A stab at perturbation theory	71

4 From black holes to stars and the Universe at large	73
4.1 The Schwarzschild solution	73
4.2 Relativistic fluids	75
4.3 How to build a star	77
4.4 The Newtonian limit	78
4.5 Modelling the Universe	82
4.6 Was Einstein right?	85
5 Binary inspiral	90
5.1 Basic celestial mechanics	90
5.2 Circular orbits	95
5.3 The Binary Pulsar	98
5.4 Eccentric orbits	99
5.5 The orbital evolution	102
6 Spinning stars and cosmic recycling	105
6.1 Rotating deformed stars	105
6.2 The Crab Pulsar	110
6.3 Contact binaries	112
6.4 Cosmic recycling	116
6.5 Spin–orbit evolution	119
7 Catching the wave	125
7.1 Resonant mass detectors	126
7.2 Gravitational waves and light beams	128
7.3 Advanced interferometers	133
7.4 An international network	137
7.5 The antenna pattern	140
7.6 The road to the future	142
7.7 Doppler tracking	148
7.8 Pulsar timing arrays	149
8 Mining the data	150
8.1 Random noise	151
8.2 Matched filtering and the optimal signal-to-noise ratio	153
8.3 Applications of matched filtering	157
8.4 Bursts searches	161
8.5 Stochastic backgrounds	163
8.6 Avoiding false alarms	165
8.7 Bayesian inference	167
8.8 Geometry in signal analysis	171

Part 2 The dark side of the universe

9 The stellar graveyard	177
9.1 White dwarfs	178
9.2 The Fermi gas model	180
9.3 Chandrasekhar's limit	182
9.4 Neutron stars	184
9.5 The rebirth of relativity	189
9.6 Weighing black holes	191
9.7 The formation of compact binaries	195
9.8 Estimating merger rates	199
9.9 Active galaxies	202
9.10 A giant at the centre of the Milky Way	204
10 Testing relativity	207
10.1 Geodesics	208
10.2 The gravitational redshift	210
10.3 Flying clocks	211
10.4 Light bending	213
10.5 Shapiro time delay	215
10.6 Light rays and black holes	216
10.7 The motion of massive bodies	218
10.8 Perihelion precession	220
10.9 The Double pulsar	221
10.10 Radial infall	223
10.11 A bit more celestial mechanics	225
11 Beyond Newton	229
11.1 Near and far-zone solutions	230
11.2 A slight aside: symmetric trace-free (STF) tensors	235
11.3 The relaxed Einstein equations	237
11.4 Iterative schemes	240
11.5 Inspiralling binaries	242
11.6 The effective one body approach	247
12 Towards the extreme	250
12.1 Matter at supranuclear densities	250
12.2 A simple model for npe matter	252
12.3 Determining the equation of state	254
12.4 Observational constraints	259
12.5 The slow-rotation approximation	260
12.6 The virial theorem	262

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12.7	The Kepler limit	266
12.8	Rotating relativistic stars	268
12.9	The quasiradial instability	272
12.10	Superfluids and glitches	274
13	From oscillations to instabilities	282
13.1	The fundamental f-mode	282
13.2	General non-rotating stars: p/g-modes	287
13.3	Calculating stellar oscillation modes	292
13.4	The r-modes	295
13.5	Gravitational-wave emission	298
13.6	What do we learn from the ellipsoids?	299
13.7	Lagrangian perturbation theory for rotating stars	305
13.8	The CFS instability	309
14	Building mountains	312
14.1	The crust	312
14.2	Energetics	316
14.3	Modelling elastic deformations	320
14.4	Searches for known pulsars	327
14.5	All-sky searches	329
14.6	The magnetic field	333
14.7	The birth of a magnetar	337
14.8	Modelling accretion	339
14.9	The low-mass X-ray binaries	344
14.10	Magnetic field burial and confinement	348
14.11	Persistent sources	351
14.12	Free precession	353
14.13	Evolution of the wobble angle	357
15	The r-mode instability	361
15.1	The instability window	362
15.2	Complicating factors	367
15.3	A simple spin-evolution model	372
15.4	Nonlinear saturation	377
15.5	Are the gravitational waves detectable?	381
15.6	Astrophysical constraints for young neutron stars	383
15.7	r-modes in accreting systems	387
16	Black-hole dynamics	391
16.1	Issues of stability	391
16.2	Scalar field dynamics	392

16.3	Gravitational perturbations	400
16.4	Quasinormal modes	405
16.5	Test particle motion	407
16.6	Taking the plunge	410
16.7	The self-force problem	412
17	Spinning black holes	418
17.1	The Kerr solution	418
17.2	Inertial framedragging	419
17.3	Kerr geodesics	421
17.4	The Newman–Penrose formalism	428
17.5	The Teukolsky equation	434
17.6	Kerr quasinormal modes	439
17.7	GW150914: A faint fingerprint	440
18	Relativistic asteroseismology	443
18.1	Relativistic fluid perturbations	443
18.2	f- and p-modes in relativity	447
18.3	The inverse problem	451
18.4	The w-modes	454
18.5	The evolving spectrum of adolescent neutron stars	457
18.6	Magnetar seismology	461
18.7	The relativistic r-modes	466
18.8	The unstable f-modes	470
19	Colliding black holes	479
19.1	The 3+1 decomposition	482
19.2	Evolving the spacetime	484
19.3	Initial data	486
19.4	Slicing conditions	489
19.5	Wave extraction	491
19.6	2 + 2 and the Bondi news	493
19.7	Milestones and breakthroughs	496
19.8	Recoil and kicks	502
20	Cosmic fireworks	508
20.1	Simulating fluids	508
20.2	The bar-mode instability	513
20.3	Tidal disruption	516
20.4	Black hole–neutron star mergers	519
20.5	Magnetohydrodynamics	522
20.6	The magnetorotational instability	525

20.7	Gravitational collapse	528
20.8	Supernova core collapse	531
20.9	Hypernovae	539
21	Anatomy of a merger	541
21.1	GW170817	541
21.2	Tidal deformation	543
21.3	The relativistic Love number	551
21.4	Dynamical tides: resonances	556
21.5	Shattering the crust	563
21.6	Merger dynamics	565
21.7	Gamma-ray bursts	572
21.8	The signature of a kilonova	578
22	Whispers from the Big Bang	581
22.1	The standard model of cosmology	583
22.2	The cosmological redshift	587
22.3	Scaling the distance ladder	589
22.4	Standard sirens	591
22.5	Geometrical optics and lensing	594
22.6	Astrophysical backgrounds	598
22.7	Pulsar timing arrays	602
22.8	AC/DC	609
22.9	Astrometry	609
22.10	Detecting a primordial background	611
22.11	Parametric amplification of quantum fluctuations	613
22.12	Phase transitions	616
22.13	Cosmic strings	617
22.14	E/B-modes	619
22.15	Twenty-nine decades of frequency	620
	Apologies and thanks	625
	References	627
	Index	661

A long time ago, in a galaxy far away, the two black holes edged closer. Dancing around each other in a nearly perfect circle. Drawn together by gravity, through the emission of gravitational waves. Faint ripples encoded the change in gravity over eons. In the last few moments the motion grew frantic. A storm of warped space and time raged as the two objects came together. An energy equal to the obliteration of several suns was released in a fraction of a second. Then it was over. All that remained was a single black hole. And empty space.

The signal moved unchanged over the vast distances of space until, after more than a billion years, it reached the Earth. When the signal was created, this insignificant blue planet hosted single cell organisms. When the signal arrived, there was an advanced civilization. A civilization curious about the Universe. A civilization with technology to catch the elusive spacetime whisper. Their advanced detectors registered a disturbance.

This was the beginning.

Opening the window

1.1 The beginning

The first direct detection of gravitational waves was announced to the world on the 11th of February 2016 with a triumphant ‘We did it!’. The signal, which had been picked up by the two LIGO detectors on the 14th of September 2015, matched the predictions from numerical simulations of the merger of a pair of black holes with masses $36M_{\odot}$ and $29M_{\odot}$, forming a larger black hole with mass $62M_{\odot}$ (Abbott *et al.*, 2016*b*). The missing mass—the equivalent of about 3 solar masses—had been radiated as gravitational waves. This extraordinary event, which only lasted a fraction of a second, was the most powerful astronomical event ever observed. It was the beginning of a new kind of astronomy.

The breakthrough detection came nearly a century after Einstein’s prediction that changes in gravity should propagate as waves (Einstein, 1916). It was an extraordinary moment of success, following decades of technology development, political wrangling to secure funding, and several false starts. It was a moment of glory, rewarding an enormous amount of patient and hard work from a lot of people.

The LIGO project was initiated in the early 1990s Abramovici *et al.* (1992) and the first generation of kilometre-scale gravitational-wave interferometers reached their initial design sensitivity in a broad frequency window in November 2005 (during the fifth science run, S5). More than one year’s worth of quality data was taken during the following science run (S6) in 2009–10. Many research papers were written, but no signals were found. After a couple of years’ downtime to improve the technology, the first ‘observing run’ (O1) of the advanced interferometers started in September 2015. The immediate detection of the black-hole signal led to a collective sigh of relief. It had been a long journey.

The first detection brought the promise of gravitational-wave astronomy into sharp focus. It was much more than a confirmation that gravitational waves exist and that we can catch them. We learned that there are double black-hole systems in the Universe and that they merge due to the emission of gravitational radiation. The observed signal agreed with the predictions from general relativity, showing the expected inspiral, merger, and ringdown phases seen in numerical supercomputer simulations (Chapter 19). It was the first test of Einstein’s theory in a dynamical, strong-field setting. The signal allowed us to identify more massive black holes than so far found in X-ray binaries, and it also provided interesting constraints on the spin of the individual black holes.

2 Opening the window

The underlying theory may be complex, but the observed signal was simple. It swept upwards in amplitude and frequency from 30 to 250 Hz in a perfect example of the anticipated chirp (see the time-frequency plots in the lower panels of Figure 1.1). At its peak, the gravitational-wave strain, $h \approx 10^{-21}$, corresponded to a luminosity equivalent to emitting the mass-energy of about 200 suns in a second. The event took place 1.3 billion light years from the Earth (Abbott *et al.*, 2016b). In terms of the Universe, it was ancient history.

Binary signals, like GW150914, carry unique information on the masses and spins of the sources. In the case of neutron stars, the gravitational waves also encode the internal structure, which depends on the state of matter at extreme densities. In essence, gravitational-wave observations have the potential to probe many fundamental physics issues. Given the weakly interacting nature of gravitational waves, the information they carry provides an important complement to electromagnetic observations. In fact, they

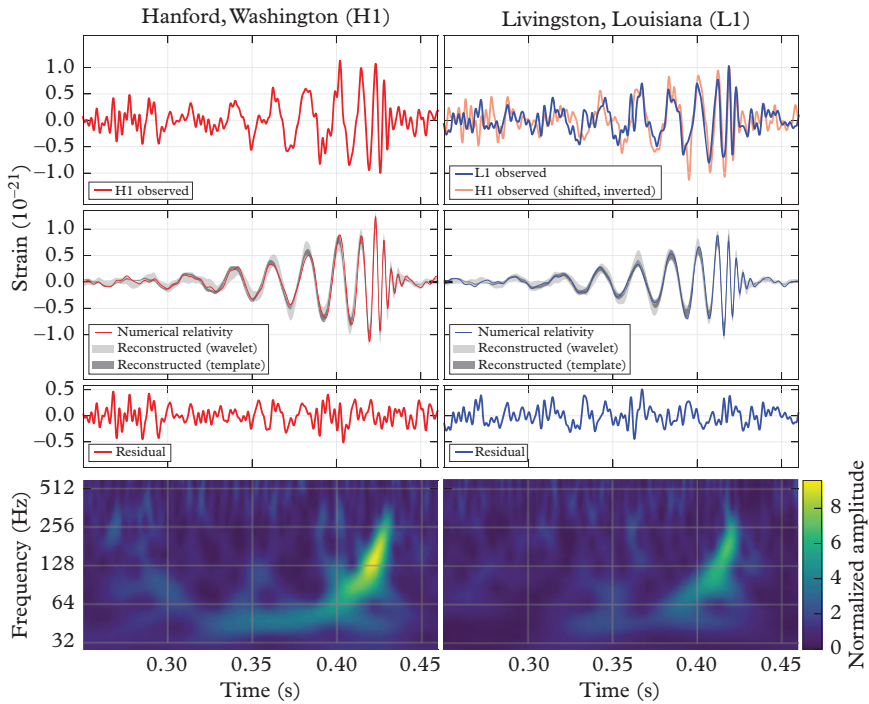


Figure 1.1 The first gravitational-wave signal (GW150914) observed by the LIGO Hanford (H1, left) and Livingston (L1, right) interferometers. The top row shows how the gravitational-wave strain varied with time in the two detectors (with a direct comparison after a time shift of 10 ms corresponding to the travel time—at the speed of light—between the two instruments). The middle row compares the signal to results from numerical relativity simulations, showing inspiral, merger, and ringdown of two coalescing black holes. The bottom row gives a time-frequency representation of the gravitational-wave strain, again showing the signal frequency and strength increasing with time. (Reproduced from Abbott *et al.* (2016b), Creative Commons Attribution 3.0 License.)

shed light on aspects that cannot be probed by traditional means, like the internal dynamics of a supernova explosion or quantum fluctuations in the very early Universe just after the Big Bang. In order to understand the wide range of possibilities, we need to explore the mechanisms that generate gravitational waves in the first place. We need to be able to predict the character of the signals and consider the challenges associated with detecting them. As this involves many complex questions, and it is important to appreciate the context, we need to start from the beginning.

1.2 A new kind of astronomy

With his theory of general relativity, Einstein revolutionized our view of space and time (Einstein, 1915). By explaining gravity in terms of the geometry of a combined spacetime he provided a fresh perspective on the Universe. This led to the introduction of exciting concepts that have become part of mainstream culture. Most notably, *black holes*, formed when massive stars die, and the *Big Bang*, the explosion which gave birth to the Universe some 14 billion years ago. Moreover, Einstein's general relativity is a *dynamic* theory of gravity, where space and time are flexible concepts. The theory predicts that changes in gravity propagate as waves, ripples in spacetime moving at the speed of light. These *gravitational waves* are elusive. For decades they caused debate and controversy¹ and, until recently, attempts to detect them proved futile.

It is not really surprising that the detection of gravitational waves proved such a challenge. Early generations of instruments may have been remarkably sensitive—from an everyday life point of view—but they would still only have been able to catch unique events in our own Galaxy and its immediate neighbourhood and such events are rare. Take supernova explosions, which occur only a few times per century in a typical galaxy, as an example. Population modelling and our understanding of stellar evolution tell us that we need to reach further out into the Universe if we want to detect such events. Exactly how far, we do not know at this point. It is relatively easy to work out the energy that must be released in order for a given source to be detectable, but very difficult to provide a reliable model of the complex physics associated with most gravitational-wave scenarios. Yet, it is clear that we will always be dealing with faint signals. This is in sharp contrast with mainstream astronomy, where observations are traditionally made at large signal-to-noise ratios.

As the sensitivity of the available detectors improved—gradually—we learned valuable lessons. It is fairly easy to identify ‘milestone’ results leading up to the breakthrough in 2015. For example, the initial LIGO–Virgo detectors were sensitive enough that they would have been able to catch a gravitational-wave burst from a Milky Way supernova, should one have occurred during the series of science runs (Abadie *et al.*, 2012). The absence of detections hardly challenged our view of the Universe, but it was nevertheless an important step. The fact that the gravitational-wave contribution to the spin-down of

¹ A meeting at Chapel Hill in January 1957 is often seen as the turning point. In particular, Richard Feynman famously provided a ‘sticky bead’ argument to demonstrate that gravitational waves must carry energy.

the Crab Pulsar—a neutron star born in a supernova recorded by Chinese astronomers in 1054—can be constrained to be less than a fraction of a percent of the observed rate (Abbott *et al.*, 2008a) may only be mildly interesting from the astrophysics point of view, but it was nevertheless a milestone achievement as it constrained the asymmetry of a distant astronomical object in a way that could not be done by other means.

Gravitational-wave astronomy is a fascinating area that involves a range of complex issues, from the development of detector technology to data-handling techniques and theory modelling. In order to progress, we need to improve on all these aspects. As we celebrate the first successful detections, it is useful to keep in mind the effort behind the success. Over decades, generations of scientists turned an impressive engineering project into an astronomical observatory. This was a spectacular achievement, but we are far from done. Future observing runs will probe a much larger volume of space. We will have more, better quality, data. Conservative population synthesis models suggest that we will detect many inspiralling compact binaries (consisting of black holes and/or neutron stars) every year. Given that such ‘bread and butter’ binary signals are well understood (and depend very little on the composition of the binary companions) and the data analysis algorithms are (more or less) developed, this should allow us to probe the parameters of such systems, shedding light on the cosmic compact binary population and the relevant formation channels.

The wider range of gravitational-wave sources put more emphasis on the involved physics and high-quality modelling of relevant astrophysical scenarios. Inevitably, this requires an exchange of expertise with mainstream astronomers. For a long time the emphasis was on detector development and data analysis strategies. As we establish this new area of astronomy, we need rapid change. We need to address challenging modelling problems. Many relevant gravitational-wave scenarios involve extreme physics that cannot be tested in the laboratory and precision searches require an understanding beyond ‘order of magnitude’ precision.

The future is, of course, bright. Once third-generation detectors, like the Einstein Telescope (Punturo *et al.*, 2010; Sathyaprakash *et al.*, 2012) or the Cosmic Explorer (Abbott *et al.*, 2017c), come on-line we will firmly be in the era of gravitational-wave astronomy. These instruments will improve the broadband sensitivity by another order of magnitude, reaching another factor of 1,000 in volume of space. This may seem remote, given that such detectors are still at the design stage, but we need to consider their promise now. We are talking about ‘big science’ and we need to understand its potential in order to argue the case for building such hugely expensive instruments. It is relevant to ask what we can hope to achieve with an Einstein Telescope, but not (necessarily) with Advanced LIGO. How much better can we do with (roughly) an order of magnitude improvement in sensitivity? Are there situations where this improvement is needed to see the signals in the first place, or is it a matter of doing better astrophysics by getting improved statistics and more precise parameter extraction? There are many interesting and complicated issues to consider.

Perhaps in contrast, it is straightforward to argue the case for a space-based detector, like the LISA project which is expected to launch in 2034 to address the European

Space Agency's science theme of the Gravitational Universe (Amaro-Seoane *et al.*, 2017). Sensitive to low-frequency gravitational waves, LISA is perfectly tuned to typical astronomical timescales (hours to minutes). If the instrument works as planned—and there is no reason to think that it should not, given the impressive results from the LISA Pathfinder (Armano *et al.*, 2018)—detection is guaranteed. In fact, many known binary systems can be used to verify that the detector is working as intended. The challenges that the LISA project faces are different. Given the number of, in principle, detectable binaries in the Galaxy, the data analyst may suffer an embarrassment of riches. The science may (to some extent) be confusion limited. However, the fact that LISA is sensitive to signals from supermassive black holes (either merging or capturing smaller objects) throughout the Universe makes it an extremely exciting mission.

On a timescale of 20 years or so we should have a network of high-precision instruments searching the skies for gravitational-wave signals over a range of up to eight decades in frequency; see Figure 1.2. These detectors will provide us with unprecedented insights into the dark side of the Universe, and allow us to probe much exciting physics. Further improvements in data quality may allow us to extract the gravitational-wave component in the cosmic microwave background. In addition, ultra-low-frequency gravitational waves are likely to have been detected by pulsar timing arrays. In parallel, we can expect to see breakthroughs in related areas of physics. Following the detection of the Higgs boson by the Large Hadron Collider, the colliders probe higher energies and may eventually find evidence for supersymmetry. Experiments aimed at detecting dark

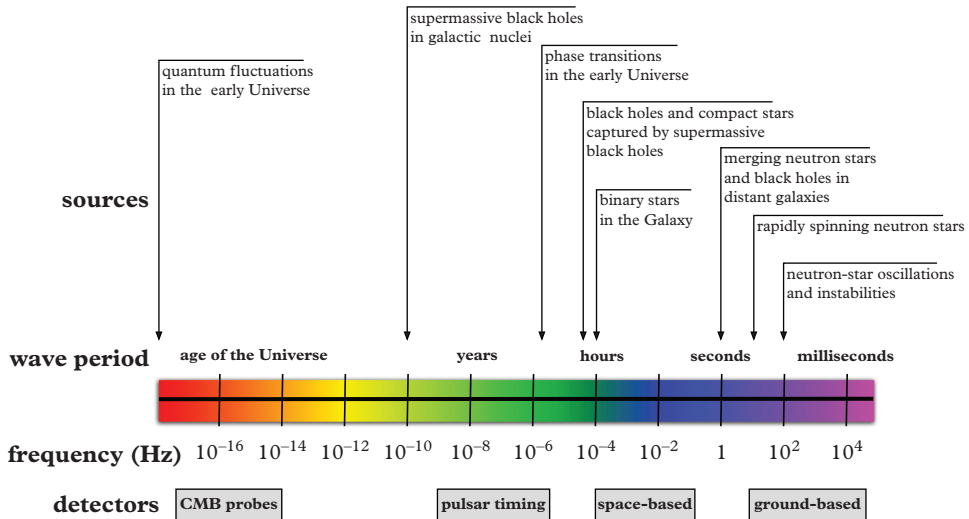


Figure 1.2 The spectrum of anticipated gravitational-wave sources and the different methods that may be used to detect them, across more than 20 decades in frequency. The physical timescales range from the age of the Universe to a fraction of a millisecond.

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Exploring the Variety of Random Documents with Different Content

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Deceivers. 1915.
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Desperation. 1917.
Discontent. 1916.
Duchess. 1915.
Dupe. 1916.
Empty Gun. 1917.
End of the Run. 1917.
Extravagance. 1915.
Fair God of Sun Island. 1915.
Faith of Her Fathers. 1915.
False Part. 1916.
Family Secret. 1916.
Fiery Introduction. 1915.
Flight of a Night Bird. 1915.
For Lack of Evidence. 1917.
Forest Nymph. 1917.
44-Caliber Mystery. 1917.
Golden Bullet. 1917.
Great Ruby Mystery. 1915.
Great Torpedo Secret. 1917.
Grip of Love. 1917.
Hair-Trigger Burk. 1917.
Heart of Gold. 1917.
Her Prey. 1915.
Homage. 1915.
Honor of an Outlaw. 1917.
Human Menace. 1915.
Husks of Love. 1916.
Idols of Clay. 1915.
In Search of a Wife. 1915.
In the Heart of New York. 1916.
Indian's Lament. 1917.
Jackals of a Great City. 1916.
June Madness. 1917.

Kiss of Dishonor. 1915.
Lady Raffles Returns. 1916.
Limb of Satan. 1917.
Lord John's Journal. (Serial)
Love's Masquerade. 1916.
Madcap Queen of Crona. 1916.
Magpie. 1917.
Manna. 1915.
Mark of a Gentleman. 1916.
Mary from America. 1917.
Masked Woman. 1916.
Master Spy. 1917.
Matty's Decision. 1915.
Measure of Leon Dubray. 1915.
Melody of Love. 1916.
Mysterious Contragrav. 1915.
Mysterious Iron Ring. 1917.
Mystery of My Lady's Boudoir. 1916.
Nature Incorporated. 1916.
Ninth Day. 1917.
Old Soldier's Romance. 1916.
On the Level. 1915.
Onda of the Orient. 1916.
Other Half. 1916.
People of the Pit. 1915.
Perilous Leap. 1917.
Phantom Fortune. 1915.
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Queen of Hearts. 1915.
Raid. 1917.
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Right-of-Way Casey. 1917.
Rose Colored Scarf. 1916.
Shattered Memories. 1915.

Should She Have Told? 1916.
Silent Man of Timber Gulch. 1916.
Six-Shooter Justice. 1917.
Soldier of the Legion. 1917.
Son of Neptune. 1916.
Soul's United. 1917.
Splash of Local Color. 1916.
Startling Climax. 1917.
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Storm Woman. 1917.
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Winning Pair. 1917.
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Woman Who Followed Me. 1916.
Won by Grit. 1917.
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Yust from Sweden. 1916.

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Going Some. 1920.
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Pest. 1919.
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*Raffles. 1930.
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*Rescue. 1929.
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*Thief in Paradise. 1924.
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Truth. 1920.
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*Blind Bargain. 1922.
*Blooming Angel. 1919.
*Bondage of Barbara. 1919.
*Bonds of Love. 1919.
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Brand. 1919.
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*Broken Chains. 1922.
*Brothers Under the Skin. 1922.
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*Christian. 1923.
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*City of Comrades. 1919.
*Come On Over. 1922.
*Concert. 1921.
*Counter Plot. 1920.
Crimson Gardenia. 1919.
Cup of Fury. 1919.
*Cupid, the Cowpuncher. 1920.
*Danger Game. 1918.
*Dangerous Curve Ahead. 1921.
Dangerous Days. 1920.
*Darn That Stocking! 1920.
*Daughter of Mine. 1919.
*Daydreams. 1918.
*Day of Faith. 1923.
*Dodging a Million. 1918.
*Dollars and Sense. 1920.
Don't Neglect Your Wife. 1921.

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*Fair Pretender. 1918.
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*Gid-ap Napoleon. 1921.
*Gimme. 1923.
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*Girl with the Jazz Heart. 1920.
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*Godless Men. 1920.
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*Grim Comedian. 1921.

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- *Hidden Fires. 1918.
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- *Hold Your Horses. 1920.
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- *In the Palace of the King. 1923.
- *Indigo Sunday. 1921.
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- *Jes' Call Me Jim. 1920.
- *Jinx. 1919.
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- *Jubilo. 1919.
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Night Rose. 1921.
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*Three Wise Fools. 1923.
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*Unwilling Hero. 1921.
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*Venus Model. 1918.
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*Wallflower. 1922.
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